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Gravitational redshift and asymmetric redshift-space distortions for stacked clusters

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ABSTRACT

We derive the expression for the observed redshift in the weak field limit in the observer’s past light cone, including all relativistic terms up to second order in velocity. We then apply it to compute the cluster–galaxy cross-correlation functions (CGCF) using N -body simulations. The CGCF is asymmetric along the line of sight owing to the presence of the small second-order terms such as the gravitational redshift (GRedshift). We identify two systematics in the modelling of the GRedshift signal in stacked clusters. First, it is affected by the morphology of dark matter haloes and the large-scale cosmic-web. The non-spherical distribution of galaxies around the central halo and the presence of neighbouring clusters systematically reduce the GRedshift signal. This bias is approximately 20 per cent for $M_{\min} \simeq 10^{14} M_{\odot} h^{-1}$, and is more than 50 per cent for haloes with $M_{\min} \simeq 2 \times 10^{13} M_{\odot} h^{-1}$ at $r > 4 \text{ Mpc } h^{-1}$. Secondly, the best-fitting GRedshift profiles as well as the profiles of all other relativistic terms are found to be significantly different in velocity space compared to their real space versions. We find that the relativistic Doppler redshift effect, like other second-order effects, is subdominant to the GRedshift signal. We discuss some subtleties relating to these effects in velocity space. We also find that the S/N of the GRedshift signal increases with decreasing halo mass.

Key words: gravitation – methods: analytical – methods: numerical – galaxies: clusters: general – large-scale structure of Universe.

1 INTRODUCTION

In general relativity, photons receive a gravitational redshift (GRedshift) when climbing out of potential wells. In the weak field limit, the magnitude of the redshift is proportional to the depth of the Newtonian potential Φ . Photons from central galaxies sitting at the bottom of the potential well of galaxy clusters are expected to be gravitationally redshifted by a larger amount than satellites and other neighbouring galaxies. The difference of the GRedshift signal with respect to the cluster centre is of the order of 10 km s^{-1} . It can, in principle, be detected by stacking a large sample of clusters. This has been predicted by Nottale (1990), Cappi (1995) and Kim & Croft (2004), and the first few tentative measurements from stacked clusters from SDSS data sets have been reported (Wojtak, Hansen & Hjorth 2011; Jimeno et al. 2015; Sadeh, Feng & Lahav 2015).

In observations, the GRedshift signal extracted from stacked clusters is related to the distortion of the cluster–galaxy cross-correlation function (CGCF), or ξ_{cg} , which originates from the distortions of the observed redshifts of galaxies with respect to the cluster centre

(which may be the centroid of the galaxies or may be taken to be the brightest cluster galaxy, BCG). In theory, ignoring the evolution of cosmic potentials and observational systematics, the observed redshift consists of five components: (1) the cosmological redshift (2), the first-order Doppler redshift from the peculiar velocity of the galaxy, (3) second-order special relativistic corrections from the peculiar velocity, (4) the peculiar GRedshift, and (5) effects associated with the fact that we observe galaxies on our past light cone. The effects of (1) and (2) result in an observed CGCF that should be front–back symmetric, while asymmetry of the CGCF along the line of sight (LOS) will arise due to the presence of (3), (4) and (5). The main goal of this study is to explore these effects on the CGCF and disentangle the GRedshift effect from them. There is also an additional effect, (6); the peculiar velocity of galaxies affects their surface brightness via beaming. Coupled to any surface brightness-dependent selection (such as an apparent magnitude limit), this results in a bias of the redshift distribution of the selected galaxies at the same order of magnitude. This last effect, unlike the others, is highly dependent on details of the luminosity function of the galaxies and how they are selected in the surveys. Here we shall focus only on those effects that are independent of how galaxies are selected.

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On large scales, the relativistic corrections to the galaxy correlation function and the resulting asymmetry of the cross-correlation function between two different ‘tracer’ populations, in our case, clusters and galaxies, has been studied in Yoo, Fitzpatrick & Zaldarriaga (2009), McDonald (2009), Challinor & Lewis (2011), Bonvin & Durrer (2011), Yoo et al. (2012), Croft (2013) and in Bonvin, Hui & Gaztañaga (2014), where some other effects such as density evolution, integrated Sachs–Wolfe effect and lensing are included. Our study will focus on the CGCF at around the scale of clusters and up to tens of $\text{Mpc } h^{-1}$. This is the (quasi-) non-linear regime where some of the theoretical predictions based on perturbative methods will break down. It is therefore necessary to employ N -body simulations for this study.

A robust detection of the GRedshift signal may provide a constraint on theories of gravity, provided that the mass and abundances of clusters are estimated self-consistently in those theories. This requires an accurate prediction of the observed redshift. Wojtak et al. (2011), for example, have modelled the effect by assuming a power-law mass function for clusters that are individually spherically symmetric and have an NFW (Navarro, Frenk & White 1996) profile, and that the observed redshifts are given as the sum of the first-order Doppler shift and the GRedshift with respect to the cluster centre. It was subsequently realized that several additional physical processes, such as the transverse Doppler redshift, the past light cone effect and relativistic beaming, would cause additional contributions that are generally of the same order of magnitude as the GRedshift signal (Kaiser 2013; Zhao, Peacock & Li 2013) and which complicate the analysis. These analyses, however, do not necessarily capture all of the relevant effects that need to be considered in order to make an accurate prediction. One shortcoming is that these analyses are not adequate to treat the ‘quasi-linear’ regime – outside the virial radius – which is observationally relevant here. Another is that, of necessity, the quantity that is measured is a *galaxy-weighted* measurement of the redshift, i.e. the mean of the GRedshift, plus other contributions, for galaxies at a given projected distance from the cluster galaxy centre. That is, it is not the simple 2-point cluster density–potential cross-correlation function, rather it is a third-order statistic $\langle n_c(0)n_g(\mathbf{r})\Phi(\mathbf{r}) \rangle / \langle n_c(0)n_g(\mathbf{r}) \rangle$, where $n_c(0)$, $n_g(\mathbf{r})$ and $\Phi(\mathbf{r})$ are the number density of central galaxies at the origin, the number density of galaxies at \mathbf{r} and the peculiar Newtonian potential at \mathbf{r} . Here we use N -body simulations to attempt to remedy these shortcomings.

The outline of the paper is as follows: In the next section, we derive an expression for the observed redshift accurate to second order in the velocities (Hubble and peculiar) and to first order in the peculiar potential and allowing for the fact that we observe galaxies on the past light cone. This provides the redshift in terms of quantities defined on a hypersurface of constant time, which is useful as the simulations provide snapshots of the galaxy positions, velocities and the peculiar gravity on such hypersurfaces.

We analyse the simulations in Section 3. This analysis reveals and quantifies two important new complicating factors. The first has to do with the fact that while, in a composite sense, clusters are spherically symmetric, individual clusters are aspherical and their surroundings are highly aspherical owing to the presence of neighbouring clusters. Coupled with the fact that the quantity one most naturally measures is the *galaxy-weighted* redshift and clumps of galaxies are correlated with potential wells, this results in a systematic bias that causes the weighted potential to increase more slowly with distance from the cluster centre than one would expect from simple models invoking an ensemble of spherical NFW profile

clusters. The second effect has to do with the fact that the galaxies are observed in velocity space rather than in real space.

2 RELATIVE REDSHIFTS ON THE PAST LIGHT CONE

We summarize the source of distortions to the observed redshifts below.

To the lowest order in peculiar velocity and potential, the distortion is associated only with the Doppler redshifts from the LOS component of the peculiar velocities. The redshift of a galaxy is

$$cz = Hx + v_x, \quad (1)$$

where x is the cosmological comoving, or conformal distance, v_x is the LOS peculiar velocity, H is the Hubble constant and c is the speed of light.

In General Relativity, GRedshift will add to the observed redshift by the amount that is proportional to the depth of the Newtonian potential Φ . The GRedshift is of the order of 10 km s^{-1} for galaxy clusters with mass $M \sim 10^{14} M_\odot h^{-1}$. In the context of special relativity, Zhao et al. (2013) realized that the transverse Doppler redshift term, dependent on v^2 , should also be added. It is guaranteed to be of the same order as the GRedshift term. Kaiser (2013) showed that there is an additional effect that is of order v^2 that comes about because the galaxies are observed on the past light cone.

In this section, we establish the connection between the Hubble and peculiar velocities of galaxies (or particles in an N -body simulation) in the vicinity of a cluster and the redshift, as would be measured by some distant observer, in the first instance, relative to the redshift of a stationary reference source that lies at the origin of coordinate system. We then generalize this to give the redshift relative to the cluster centre.

Since these relative redshifts are very small, we may analyse this using Newtonian gravity with gravitationally induced wavelength shifts $\delta\lambda/\lambda_{\text{em}} = -\Phi/c^2$, where $\delta\lambda \equiv \lambda_{\text{obs}} - \lambda_{\text{em}}$, and using special relativity to compute the Doppler shifts. Since these wavelength shifts are multiplicative, we can simply deal with this as treating their logarithms as additive. Furthermore, since the (total) potential is of the order of the square of the total velocity (i.e. Hubble plus peculiar), it is sufficient here to work to second order in velocities and first order in the potential. Also, since the potential will generally be evolving on the dynamical time-scale and the velocities are highly non-relativistic, one may ignore the evolution of the potential in the relatively tiny light travel time.

More specifically, the relative redshift may be calculated in terms of peculiar velocities and the peculiar potential – which is the solution of Poisson’s equation with the density perturbation as the source term and which is what appears in the equations of motion that are solved in N -body simulations – as follows: First, we may calculate the relative redshifts that would be observed in the fictitious situation where the density of the universe is unperturbed but where one is observing a set of particles that have peculiar velocities. This is given by (one plus) the cosmological redshift – which is just inversely proportional to the scalefactor at the time of emission – multiplied by the relativistic Doppler shift for a moving particle with respect to a comoving particle (i.e. one with vanishing peculiar velocity). Note that the peculiar velocity here is the peculiar velocity at the emission time, whereas what we are supplied with most conveniently is the output of N -body simulations on a hypersurface of constant proper time. So it is necessary to allow for the ‘Hubble drag’ that causes a change of peculiar velocity with time. If we now ‘switch on’ the effect of gravity, we need to include the lowest order

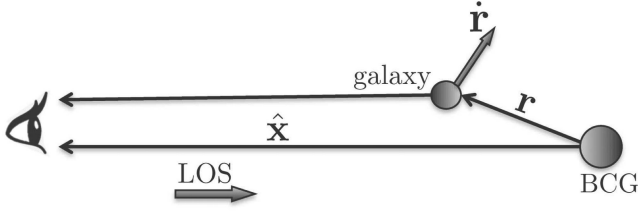


Figure 1. Illustration of the BCG–galaxy system. The observer is on the left, assumed to be static in conformal coordinates with respect to the BCG. \mathbf{r} is the conformal position of the galaxy with respect to the cluster centre, and $\dot{\mathbf{r}}$ is its conformal velocity. Photons received by the observer at the conformal time η_0 are emitted at a different conformal time from the galaxy and the BCG. During this interval of look-back time, the Universe expands, the galaxy moves and may be accelerated with respect to the BCG. These give rise to the second-order terms in equation (17).

GRedshift by adding the appropriate Newtonian peculiar potential to the fractional frequency shifts and we need to allow for the fact that there is not just Hubble drag but also peculiar gravitational acceleration that changes the peculiar velocity.

Let us suppose we are given a set of galaxies coordinates and velocities (or those of the particles in a N -body simulation) on a constant proper-time hypersurface. More specifically, let us assume that for each galaxy, we have the position \mathbf{r} , this being the comoving coordinate times the scalefactor a ; the peculiar velocity $\mathbf{v} = d\mathbf{r}/dt$ (from which we can obtain the conformal velocity $\dot{\mathbf{r}} \equiv d\mathbf{r}/d\eta = a\mathbf{v}$), where conformal time is defined, up to a constant, by $d\eta = dt/a(t)$.

Let us also assume that we are provided with the peculiar potential Φ and its gradient $\mathbf{g} = -\nabla_{\mathbf{r}}\Phi$, again at some given conformal time $\eta = \eta_0$.

We will use units such that $c = 1$ temporally, and put back c for the final expression of our derivation. If we set $a = 1$ at the output time, then \mathbf{v} and $\dot{\mathbf{r}}$ are identical at that time and separations in \mathbf{r} are proper separations in physical units.

Extending off the time-slice $\eta = \eta_0$, a galaxy will have trajectory

$$\mathbf{r}(\eta) = \mathbf{r} + (\eta - \eta_0)\dot{\mathbf{r}} + \dots, \quad (2)$$

where \mathbf{r} , and $\dot{\mathbf{r}}$, without an argument, indicate the values at η_0 and \dots indicates terms that are of second or higher order in conformal look-back time $\Delta\eta \equiv \eta - \eta_0$.

As illustrated in Fig. 1, we will place the observer event at some large distance along the (minus) x -axis, and at the time such that the observer receives photons that left the origin (which we will ultimately take to be the centre of the cluster) at time η_0 . The equation of the surface in \mathbf{r} -space that contains the points on the observer's past light cone with conformal time η is

$$\eta = \eta_0 - \hat{\mathbf{x}} \cdot \mathbf{r}(\eta) + \dots, \quad (3)$$

where $\hat{\mathbf{x}}$ is the unit vector parallel to the x -axis and where we are ignoring the fact that the coordinate speed of light is not exactly unity because of the metric perturbations (this introduces errors of order $v \times \Phi$, which we may safely neglect). This formula gives the conformal emission time of a photon from a particle at relative position \mathbf{r} that is received at the same time as a photon that leaves the origin at time η_0 .

For simplicity, here we are making the ‘plane–parallel’ approximation, which is valid for sufficiently distant clusters.

Substituting (2) in (3) yields the conformal look-back time in terms of \mathbf{r} and $\dot{\mathbf{r}}$:

$$\Delta\eta = -\hat{\mathbf{x}} \cdot \mathbf{r} / (1 + \hat{\mathbf{x}} \cdot \dot{\mathbf{r}}) = -\hat{\mathbf{x}} \cdot \mathbf{r} + (\hat{\mathbf{x}} \cdot \mathbf{r})(\hat{\mathbf{x}} \cdot \dot{\mathbf{r}}) + \dots \quad (4)$$

or, with $x = \hat{\mathbf{x}} \cdot \mathbf{r}$ and $\dot{x} = \hat{\mathbf{x}} \cdot \dot{\mathbf{r}} = v_x$:

$$\Delta\eta = -x + x\dot{x} + \dots \quad (5)$$

We may use this to calculate the (inverse) redshift associated with the expansion of the universe during this look-back interval:

$$(1+z)^{-1} = \frac{a(\eta)}{a(\eta_0)} = 1 + \frac{\dot{a}}{a}\Delta\eta + \frac{1}{2}\frac{\ddot{a}}{a}(\Delta\eta)^2 + \dots \quad (6)$$

or

$$1+z = 1 - \frac{\dot{a}}{a}\Delta\eta + \left[\left(\frac{\dot{a}}{a} \right)^2 - \frac{1}{2}\frac{\ddot{a}}{a} \right] (\Delta\eta)^2 + \dots \quad (7)$$

or with $\Delta\eta$ given by (5)

$$1+z = 1 + \frac{\dot{a}}{a}x - \frac{\dot{a}}{a}x\dot{x} + \left[\left(\frac{\dot{a}}{a} \right)^2 - \frac{1}{2}\frac{\ddot{a}}{a} \right] x^2 + \dots \quad (8)$$

This is *not* the redshift of a real galaxy with time-slice position \mathbf{r} and velocity $\dot{\mathbf{r}}$ (at the time η when it intercepts the observer's past light cone), rather it is the redshift of a stationary source that is co-located with that galaxy at that time relative to a stationary source at the origin $\mathbf{r} = 0$ in a fictitious universe with no structure and therefore no peculiar GRedshift. To obtain the redshift of the actual particle of interest, we need to multiply (8) by the appropriate Lorentz boost factor and we need to include the peculiar GRedshift.

The Doppler shift (the redshift of the emitting galaxy as seen by a co-located stationary observer) is (Einstein 1907)

$$(1+z)_{\text{Doppler}} = \frac{1+\dot{x}}{\sqrt{1-v^2}} = 1 + \dot{x} + v^2/2 + \dots, \quad (9)$$

but here \dot{x} is the peculiar velocity at the time of emission, which differs (at second order) from the velocity at the output time η_0 . The equation of motion for the peculiar velocity is

$$\dot{\mathbf{v}} = \mathbf{g} - H\mathbf{v}, \quad (10)$$

where \mathbf{g} is the peculiar acceleration and the second term is the ‘Hubble drag’ term that arises because peculiar velocities are defined to be with respect to the expanding (constant comoving coordinate) observers. Thus, the LOS velocity appearing in (9) is

$$\dot{x}(\eta) = \dot{x}(\eta_0) - (g_x - H\dot{x})x, \quad (11)$$

where we have used $\Delta t = \Delta\eta = -x$.

Multiplying (8) and (9) and keeping up to second-order terms and adding the peculiar GRedshift gives, for the redshift of the galaxy with respect to that for a stationary emitter at the origin,

$$cz = Hx + v_x + v^2/2c - \Phi/c - xg_x + Hxv_x/c + [H^2 - \ddot{a}/(2a^2)] x^2/c, \quad (12)$$

where we have put back the speed of light. The above equation fully accounts for the observed redshift relative to a stationary emitter on the past light cone to second order (if the potentials are not evolving). We call the total distortion to the Hubble term induced by all the other terms the ultimate redshift-space distortion (uRSD).

The above formula gives the redshift of a galaxy (or particle in a simulation) relative to a stationary source lying at $\mathbf{r} = 0$. More observationally relevant is the redshift relative to the centre of the cluster. This might be defined to be the BCG, or it may be defined to be the centroid of the cluster members. The above formula can be used to obtain the redshift of the BCG, and one might naively imagine that the relative redshift of the galaxy relative to the BCG would be the difference of these. But this is not the case; at least when working to second-order precision. The relevant relative redshift is

$1 + \delta z = \lambda_{\text{obs}}/\lambda'_{\text{obs}} = (1 + z)/(1 + z')$, where λ'_{obs} is the observed wavelength for light received from the centre and z' is the corresponding redshift. Because z' appears in the denominator, we cannot simply take $\delta z = z - z'$.

In order to deal with this situation, or the yet more complicated situation where the centre of the cluster is the centroid of the members, it is more convenient to work in terms of z_* , the logarithm of $1 + z$. Then the relative z_* between the galaxy and the BCG is just the difference of z_* for these objects. Also, the z_* of the centroid is just the average of the z_* values for the cluster members. At second order,

$$z_* = z - ((Hx + v_x)/c)^2/2, \quad (13)$$

and we have

$$cz_* = Hx + v_x + v^2/2c - v_x^2/2c - \Phi/c - xg_x + [H^2 - \ddot{a}/(a^2)]x^2/(2c). \quad (14)$$

To estimate the impact of those second-order terms for real observations, there is no unique way, as it depends on what convention the ‘observed redshift’ is adopted. For example, Wojtak et al. (2011) took the LOS ‘velocity’ of a galaxy with respect to the BCG as $\Delta V_{\text{LOS}} = c \frac{z - z_c}{1 + z_c}$. In terms of z_* , it becomes

$$\frac{\Delta V_{\text{LOS}}}{c} = \frac{e^{z_*} - e^{z_{*c}}}{e^{z_{*c}}} \approx \Delta z_* + \frac{1}{2} \Delta z_*^2, \quad (15)$$

where $\Delta z_* = z_* - z_{*c}$ and the subscript c denotes quantities for the BCG. However, if one uses $(z - z_c)/(1 + z)$ instead of having $(1 + z_c)$ in the denominator, then equation (15) would become

$$\frac{\Delta V_{\text{LOS}}}{c} = 1 - \exp(-\Delta z_*) \approx \Delta z_* - \frac{1}{2} \Delta z_*^2, \quad (16)$$

which is different from equation (15) at the second order. Nevertheless, in this work, we choose the convention of equation (15) as an example for illustration. Combining equations (13)–(15), we have

$$\Delta v_{\text{LOS}} = Hx + \Delta v_x + \Delta v^2/(2c) + [(\Delta v_x)^2 - \Delta v_x^2]/(2c) - \Delta \Phi/c - xg_x + Hx\Delta v_x/c + (H^2 - \ddot{a}/a^2/2)x^2/c. \quad (17)$$

Δx and Δv_x are differences of a galaxy’s LOS distance and peculiar velocity with respect to the BCG, respectively, so $\Delta x = x$ and $\Delta v_x \neq v_x$ by definition, and $\Delta v_x = v_x - v_{xc}$, $\Delta v^2 = v^2 - v_c^2$ and $\Delta \Phi = \Phi - \Phi_c$.

The various terms in the uRSD (equation 12) can be understood as follows:

- (i) The first two terms on the right-hand side are the Doppler shift from the total (i.e. Hubble + peculiar) velocity.
- (ii) We then have the transverse Doppler effect and the peculiar GRedshift.
- (iii) Next, we have minus the product of the LOS displacement and the LOS acceleration; these tend to be anticorrelated for overdense systems and combine to give the (positive redshift) effect shown in (Kaiser 2013), but in Section 6.3, we will see that the situation is more complicated in velocity space.
- (iv) Next, we have a second-order term Hxv_x/c , which is the product of the Hubble and peculiar velocities. In the virialized region, these will be uncorrelated, but in the outskirts of a cluster, they will be anticorrelated, so should give a negative contribution to the mean redshift. Again, the situation in velocity space and further from the cluster centre may be different.
- (v) We then have the quadratic term (in x) that comes from the combination of the background GRedshift and Doppler effects (it

is present even if v and Φ are zero). In a situation where the density of galaxies is constant in real space, this will introduce, at leading order, a linear ramp in the density. However, in analyses of GRedshift such as those of Wojtak et al. (2011) and Jimeno et al. (2015), this gets removed because they fit for the local large-scale gradient using the density of galaxies well separated in velocity from the cluster. Similar effects arise from the fact that the cluster will be at finite distance, so a beam through the cluster in which the distribution of galaxies is measured will be broadening, and also because of variation of the selection function. We will assume that the process for fitting the background density ramp has removed all of these.

As mentioned, there is one final complication in that the surface brightness of a galaxy at a given distance and light emission time depends on the peculiar velocity. This couples to the selection criterion. One could deal with this using the ‘Poisson sample’ model in which we assume that galaxies in a given volume element are drawn from the luminosity function, and are then selected according to observational criteria, and where the overall normalization includes the space density of haloes as a multiplicative factor. The ‘Doppler boosting’ modulation can be incorporated by giving a weight to the haloes extracted from the simulations. We note that unlike the other effects, this is sensitive to exactly how galaxies are selected, which, in turn, is a function of distance to the cluster.

On the right-hand side of the above equation, apart from the first two terms, which give rise to the conventional RSD, all the other terms will cause asymmetry in the CGCF. The RSD signal coming from the Δv_x term is expected to be dominant over all other effects. Recovering the asymmetry signal from the observed CGCF resulting from the uRSD is nontrivial. In principle, other astrophysics processes like dust extinction can also complicate the CGCF, which is beyond the scope of our study with N -body simulations. We focus on gravity and light cone effect in this work. Our goal is to use N -body simulations to quantify each of these terms and so determine the contamination of the GRedshift signal. Before we quantify the RSD of the CGCF caused by each effects of the of equation (17), we first study the GRedshift signal using the full 3D, real space information of particles in our N -body simulations. We show that even in this ideal situation, there is a subtle systematic effect when one assumes spherical symmetry when stacking.

3 TESTING THE SPHERICAL ASSUMPTION FOR STACKED CLUSTERS

The first measurement of the GRedshift signal from stacked clusters has been conducted by Wojtak et al. (2011) using the SDSS MaxBCG group catalogue from (Hao et al. 2010). The idea is that the BCG is likely to live close to the bottom of the potential well of the host halo, while other galaxies (satellites and field galaxies) further away from the centre of the halo tend to occupy locations where the gravitational potential is shallower. Therefore, there are relative blueshifts of the spectra coming from other galaxies relative to that of the BCGs seen by the observer.

In observations, the GRedshift signal of stacked clusters arises from the galaxy number weighted gravitational potential profile,

$$\bar{\Phi}^{\text{obs}}(r) = \frac{\int (dN_c/dM)dM \int n_{c,g}(\mathbf{r})[\Phi_c - \Phi(\mathbf{r})]d\Omega}{\int (dN_c/dM)dM \int n_{c,g}(\mathbf{r})d\Omega}, \quad (18)$$

where Φ_c and $\Phi(\mathbf{r})$ are the Newtonian potentials at the centre of the cluster and at the position \mathbf{r} , dN_c/dM is the number density of clusters per unit mass (i.e. the cluster mass function) and $n_{c,g}(\mathbf{r})$ is the number (density) of observed galaxies in the cluster at \mathbf{r} . It is

important to note that \mathbf{r} is a 3D vector, and not a scalar. Previous studies in this subject usually take r as a scalar, which implicitly assumes spherical symmetry for each cluster as well as for the stacked cluster composite. It is reasonable to expect that the stacked cluster will be close to being spherically symmetric as long as the sample is large. However, individual clusters are not spherically symmetric and contain sub-structures. Hence, in the stack, their potentials will be weighted more strongly at the locations where there are more galaxies (and more mass), and since more mass is associated with deeper local potentials, this will bias the weighted potential relative to the spherically averaged potential. This is one of the key points that we aim to address. By writing down \mathbf{r} as a 3D vector, the above expression gives the mass-weighted (or galaxy number-weighted) potential profile.

To be explicit, for the mass-weighted case, the averaged potential at a given r for the stacked clusters is affected by the contribution of each cluster system in three different ways. First, the composite is weighted by the number of galaxies contributed by each cluster at each r . More massive clusters generally contribute more weight. Secondly, within each spherical shell, the potential is weighted more strongly in the directions where there is more mass and more galaxies. This subtle effect can bias the average potential. Thirdly, the potential profiles are weighted by the number of clusters. This is represented by the outer integral over the halo mass function in equation (18), where, in observations, the selection function should also be incorporated properly.

Moreover, when measuring the GRedshift profiles at a few times the virial radius of the main halo, it is oversimplistic to extrapolate an analytic halo profile such as the NFW profile (Navarro et al. 1996) to such large radii. The presence of neighbouring galaxies and clusters will generate local potential minima. They will alter the shape of the potential profile. Naively, one may expect that these are random fluctuations and they will cancel out when averaging over a large sample of clusters. However, the fact that local potential minima are correlated with the local overdensity of galaxies will mean that the galaxy-weighted potential will be biased by such fluctuations in the potential.

Equation (18) is essentially what one will measure from observations, whereas for the modelling, spherical symmetry is usually assumed for each cluster. This is equivalent to dropping the angular dependence from equation (18):

$$\bar{\Phi}^{\text{iso}}(r) = \frac{\int (dN_c/dM) n_{c,g}(r) [\Phi_c - \Phi(r)] dM}{\int (dN_c/dM) n_{c,g}(r) dM}. \quad (19)$$

In this case, the potentials are weighted equally in all directions, but at each r , we retain the relative weighting of clusters of different masses. We will refer to this case as ‘isotropic with halo mass weighting’.

An even more naive model for the potential profile is to give equal weight to each spherical shell for each cluster. In this case, equation (19) becomes

$$\bar{\Phi}^{\text{idl}}(r) = \frac{\int (dN_c/dM) [\Phi_c - \Phi(r)] dM}{\int (dN_c/dM) dM}, \quad (20)$$

which we refer to this idealized case as equal weighting. We use cosmological N -body simulations to see how, in practice, the three stacked potential profiles of equations (18)–(20) differ.

4 THE SIMULATION SET-UP

We use the Millennium simulation for our study (Springel et al. 2005). The simulation was run in the concordance Λ cold dark

matter model, with $\Omega_m = 0.25$, $\Omega_\Lambda = 0.75$, $h = 0.73$, $\sigma_8 = 0.9$ and $n = 1$. It has 2160^3 particles in a box of $500 \text{ Mpc } h^{-1}$ on a side. The particle mass is $8.6 \times 10^8 \text{ M}_\odot h^{-1}$. We focus on haloes with the mass of $M > 10^{13} \text{ M}_\odot h^{-1}$. They have at least 10^4 particles. The softening length of the force is $5 \text{ kpc } h^{-1}$. The high resolution of the simulation enables us to probe the gravitational potential profiles deep into the halo centres.

Friends-of-Friends (FoF) groups are identified in the simulation using a linking length of 0.2 times the mean inter particle separation (Davis et al. 1985). Sub-haloes are found starting from FoF groups using SUBFIND (Springel et al. 2001). To identify the halo centre, we first find the most massive sub-halo within each FoF halo, then sort all particles within the sub-halo according to their total energy. The most bound particle, i.e. the one with the lowest total energy, is treated as the halo centre. Note that this centre is usually different from the minimum of the potential of the group. Halo masses are defined as the mass around the halo centre and within the radius r_{200} , where r_{200} is the radius within which the mean density is 200 times of the critical density of the Universe.

We use the GADGET-3 (Springel et al. 2008) code to evaluate the peculiar potentials for all particles. In each cluster, we define the potential of the cluster centre to be the mean of that of particles within a core radius of $3 \text{ kpc } h^{-1}$ of the cluster centre. Changing the size of the core radius simply shifts the profiles of $\Phi_c - \Phi(r)$ up and down, but does not alter their shapes.

To estimate the stacked potential profiles for the case of isotropic averaging, we insert massless test particles uniformly distributed on spherical shells around each halo centre. We use HEALPIX (Górski et al. 2005) to generate 3072 pixel coordinates over 4π steradians. The mean spacing of pixels is about 3.6° . Along the radial direction, 20 spherical shells are distributed logarithmically per decade, starting at $0.01 \text{ Mpc } h^{-1}$ from the halo centres out to $30 \text{ Mpc } h^{-1}$. Convergence of the results has been tested in the radial and tangential directions by doubling the sampling rate along both directions, respectively, and we have found no noticeable change in our results. The $z = 0$ snapshot is used for our study.

5 RESULTS FROM 3D PROFILES

Fig. 2 gives two examples of clusters from the Millennium simulation. The matter distribution of these clusters is far from spherical, especially at large radii from the cluster centres. There are filamentary structures and in some cases, massive neighbouring clusters within a $10 \text{ Mpc } h^{-1}$ radius. The potentials projected along one major axis of the simulation box clearly illustrate the complex morphology of the potential within the $10 \text{ Mpc } h^{-1}$ radius regions. The main central haloes are associated with the deep potential wells. The bottom panels show the corresponding mass-weighted potential profiles Φ^{obs} of equation (18), which are similar to what will be observed (except that this is in real space), and the spherically averaged profiles, Φ^{iso} , evaluated using the isotropic averaging of equation (19).

The main halo shown on the left is relatively massive, $10^{15} \text{ M}_\odot h^{-1}$. The shape of the potential well may seem symmetric close to the bottom of the minimum (middle left-hand panel). However, it is not, as it can be seen that the colour distribution is clearly not symmetric. This indicates variation in the projected mass density. It is consistent with the picture in the top left-hand panel, where a bar-shaped core is visible close to the centre. Along the direction of elongation near the core of the cluster, the potential values are slightly deeper than those along the perpendicular short axis. Seen in 3D, the potential well looks like a valley floor, where more mass

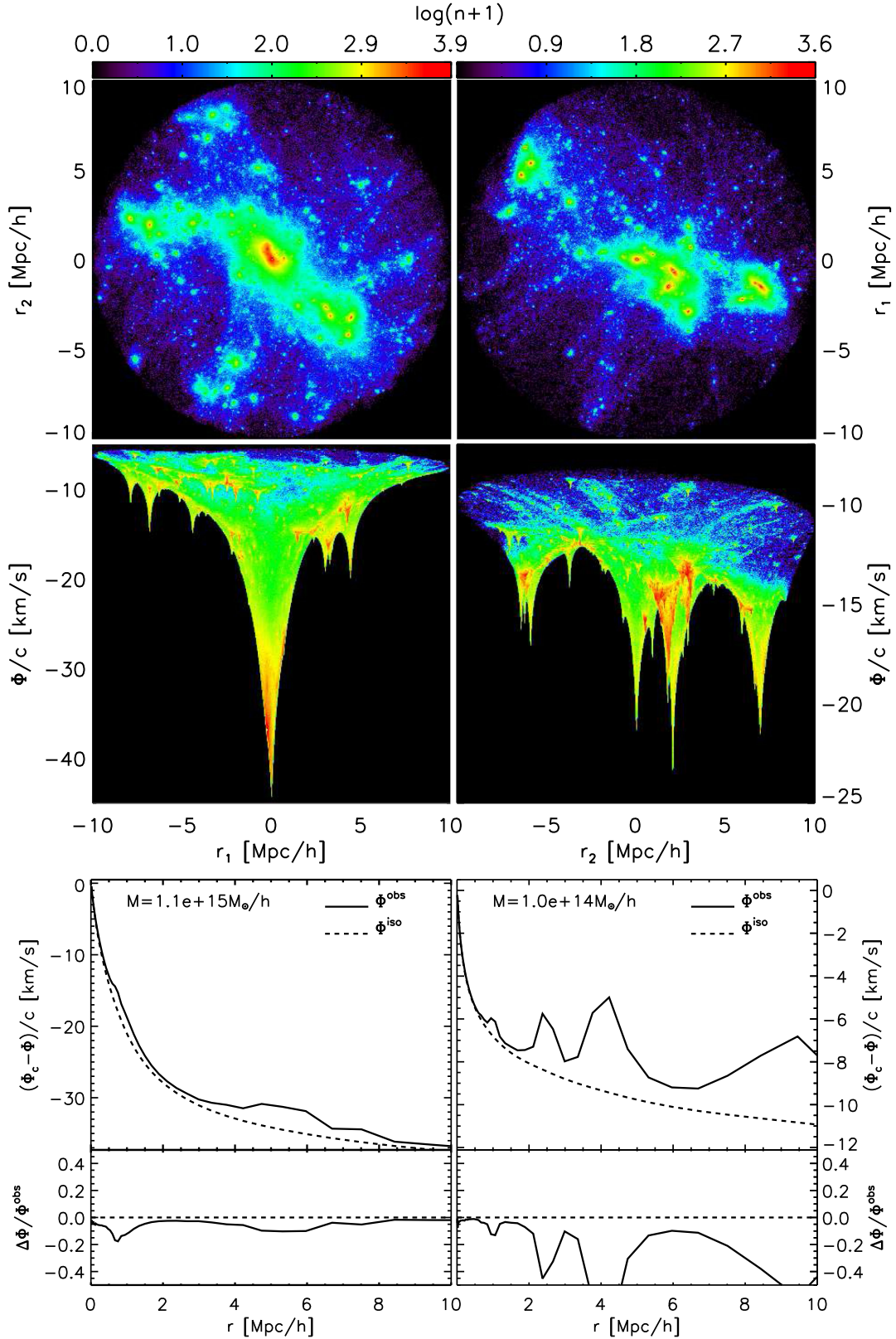


Figure 2. Top panels: particle distributions within a $10 \text{ Mpc } h^{-1}$ radius of the main halo centre projected along one major axis of the simulation box. The colour displays the number of particles in each pixel, n , as indicated by the colour bar. Middle panels: the same regions and colour coding as the top panels but now showing the value of the potential of each particle on the y-axis. Sub-haloes and neighbouring structures generate local potential minima. Bottom panels: the GRedshift profiles relative to the cluster centres. The dashed lines show the spherical averaged profile, Φ^{iso} , which is the same as obtained by isotropic weighting from the halo centres. Sub-haloes and neighbours cause the mass-weighted profiles Φ^{obs} to be biased low compared to spherical averaging. This is similar to observations where the observed profiles are weighted by galaxies.

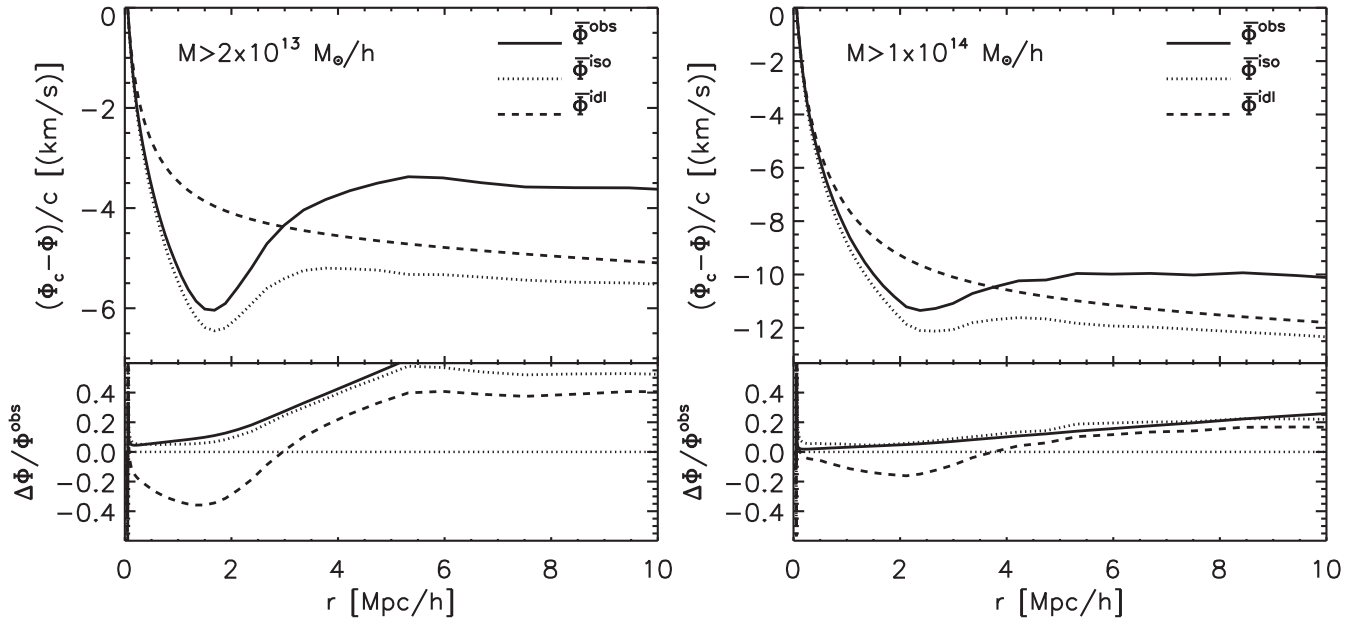


Figure 3. The gravitational potential profiles for stacked haloes in different mass bins as indicated in the legends. Potential values for halo centres, Φ_c , are approximated by averaging those of all particles within a core radius of $r_c = 3 \text{ kpc } h^{-1}$ around the location of each most bound particle. Solid lines show results from the mass-weighted $\bar{\Phi}^{\text{obs}}$ of equation (18). The dotted lines are for the isotropically weighted case, $\bar{\Phi}^{\text{iso}}$, of equation (19). The dashed lines are for the idealized case of $\bar{\Phi}^{\text{idl}}$ from equation 20, where we additionally give equal weight to each halo rather than weighting them according to the mass they contribute to the shell. Bottom panels show the fractional differences $(\bar{\Phi}^{\text{obs}} - \bar{\Phi}^{\text{iso}})/\bar{\Phi}^{\text{obs}}$ (dotted lines) and $(\bar{\Phi}^{\text{obs}} - \bar{\Phi}^{\text{idl}})/\bar{\Phi}^{\text{obs}}$ (dashed lines). The absolute differences between $\bar{\Phi}^{\text{obs}}$ and $\bar{\Phi}^{\text{iso}}$ are found to follow a simple linear function when the radius is rescaled by r_{200} , i.e. $\Delta\Phi/c \approx 0.25r/r_{200} \text{ km s}^{-1}$, as shown by the solid curves in the bottom panels.

is concentrated along the valley. At small radii, this causes the mass-weighted potential profile Φ^{obs} to be shallower than the spherically averaged profile, Φ^{iso} . One can see the corresponding kink within $1 \text{ Mpc } h^{-1}$ in the lowermost left-hand panel.

The neighbouring structures also induce local potential minima. For relatively small neighbours, the neighbouring potential minima are shallower than the central ones. This is the case for the left-hand figure. The non-spherical nature of the main halo and the presence of neighbours make the mass-weighted potential profiles Φ^{obs} shallower than the case of spherical averaging Φ^{iso} , as seen in the bottom panels.

As the mass of the main halo gets smaller, the chance of having a comparably massive neighbouring structure increases. In some cases, a neighbouring cluster can be even more massive than the main central halo, as shown by the example in the right-hand panels of Fig. 2. Here the mass of the main halo is $10^{14} M_{\odot} h^{-1}$. At about $5 \text{ Mpc } h^{-1}$ from the chosen main halo centre, one of neighbouring haloes has a deeper potential minimum than the main halo. The very massive neighbours cause strong biases of the potential profile (bottom right-hand panel). Note that the neighbouring systems shown in Fig. 2 are not sub-structures of the main halo. They are essentially correlated large-scale structures outside the virial radius of the main halo.

5.1 The stacked potential profiles

With the individual observed potential profiles Φ^{obs} always being biased low compared to the spherically averaged Φ^{iso} , it is clear that the stacked profiles cannot be given by averaging the Φ^{iso} profiles, even if the stacked cluster system is perfectly spherically symmetric. Results for two different halo mass ranges are shown in Fig. 3. As expected, the stacked profiles of $\bar{\Phi}^{\text{obs}}$ are systemati-

cally lower, resulting in less negative values of $\Phi_c - \bar{\Phi}^{\text{obs}}$ in Fig. 3 (solid lines) than the corresponding spherically averaged profiles, $\Phi_c - \bar{\Phi}^{\text{iso}}$ (dotted lines). This indicates that the blueshifts of the surrounding galaxies relative to the central BCGs will be smaller than predicted by the assumption of spherical symmetry. We find the absolute difference between $\bar{\Phi}^{\text{obs}}$ and $\bar{\Phi}^{\text{iso}}$ for $M > 10^{14} M_{\odot} h^{-1}$ can be well approximated by a linear function when the radii are rescaled by r_{200} , i.e. $\Delta\Phi/c \approx -0.25r/r_{200} \text{ km s}^{-1}$. This approximation is shown by the solid lines in the lower panels of Fig. 3. This approximated relation also holds for the case of $M > 2 \times 10^{13} M_{\odot} h^{-1}$ at $r < 5 \text{ Mpc } h^{-1}$. We find that this approximation holds for a wide range of minimum halo masses M_{min} , from 10^{13} to $10^{15} M_{\odot} h^{-1}$. In terms of fractional differences, $(\bar{\Phi}^{\text{obs}} - \bar{\Phi}^{\text{iso}})/\bar{\Phi}^{\text{obs}}$, these are stronger when M_{min} is small. For $M_{\text{min}} = 2 \times 10^{13} M_{\odot} h^{-1}$ shown in the left-hand panel of Fig. 3, the bias is approximately 60 per cent at $r > 5 \text{ Mpc } h^{-1}$. For $M_{\text{min}} = 10^{14} M_{\odot} h^{-1}$ (right-hand panel), the bias varies from a few per cent to more than 20 per cent. We also find that in an extreme case when clusters with relatively low halo mass ($1 \times 10^{13} < M < 2 \times 10^{13} M_{\odot} h^{-1}$) are used, the mass-weighted potential profiles are very close to zero at most scales due to the presence of neighbouring structures.

It is noticeable that both $\Phi_c - \bar{\Phi}^{\text{obs}}$ and $\Phi_c - \bar{\Phi}^{\text{iso}}$ have troughs at approximately $2 \text{ Mpc } h^{-1}$ caused by the presence of the secondary potential wells at radii greater than $2 \text{ Mpc } h^{-1}$, which cause $\Phi_c - \bar{\Phi}^{\text{obs}}$ to become less negative at $r > 2 \text{ Mpc } h^{-1}$. The rises on the right-hand side of the troughs seen in Fig. 3 are more pronounced for $\Phi_c - \bar{\Phi}^{\text{obs}}$ than for $\Phi_c - \bar{\Phi}^{\text{iso}}$ as mass weighting gives more weight to the secondary potential wells. The troughs are also stronger for lower values of M_{min} as the chance of having more massive neighbours is greater. Note that the troughs are not seen in previous models in the literature, e.g. Wojtak et al. (2011). One may suspect that this might be due to the fact that the profiles we show here are

from 3D averaging. We will show in the next section that even when quantified by projected distances, the troughs in the profiles remain. So projection effects are not the explanation for the absence of the troughs.

Note that the biases of the dotted lines with respect to the solid lines are purely due to the assumption of spherical symmetry. For completeness, we also compare them with the most simplistic case where the weighted contributions from each of the individual haloes at each r are assumed to be equal, i.e. Φ^{idl} of equation (20). These results are shown by the dashed lines in Fig. 3. No trough is seen and the profiles are smooth and monotonic. The biases for this case of $\bar{\Phi}^{\text{idl}}$ versus $\bar{\Phi}^{\text{obs}}$, defined as $(\bar{\Phi}^{\text{obs}} - \bar{\Phi}^{\text{idl}})/\bar{\Phi}^{\text{obs}}$, are 40 per cent (left-hand panel) and 20 per cent (right-hand panel) at their maxima, as shown by the dashed curves in the bottom panels of Fig. 3.

In summary, haloes are, in general, ellipsoidal rather than spherical. Within the virial radius of a halo, there is more mass concentrated along the long axis of the halo. The higher mass concentration generates deeper potential valleys along the major axis. The mass-weighted potentials are therefore higher than the case of spherical averaging. Outside the virial radius of the halo, the matter distribution is even further from being spherical distributed. The filamentary structures and neighbouring haloes embedded in them create deep secondary potential minima. These tend to decrease the potential difference with respect to the cluster centre. The difference for the potential profiles between mass weighting versus spherical averaging is comparable to the model differences between the predicted GRedshifts for some modified gravity theories and General Relativity (Wojtak et al. 2011). The biases, if not accounted for, may confuse the interpretation of the observed signal. However, we will show in the next section that the picture we have presented so far will change significantly when observing particles/galaxies in velocity space. Also, the predicted signal will be altered by the other terms arising from the treatment of the past light cone.

6 THE ULTIMATE REDSHIFT SPACE DISTORTION

Results from the previous sections use 3D spatial information for clusters in real space, which is useful for understanding the physics. In this section, we take one step closer to the observations by placing the simulated clusters in velocity space and including the past light cone effects. We will quantify each term on the right-hand side of equation (17) and compare them to the GRedshift signal. We summarize the key steps for doing this with simulations:

(i) Each particle including the one representing the BCG around a cluster is assigned an observed LOS velocity with respect to that of the BCG according to equation (17). This, in essence, turns the comoving coordinates of particles (galaxies) into observed redshifts in the past light cone, which include the Hubble expansion, peculiar velocities, GRedshifts and all the other second-order effects specified in the equation.

(ii) We then measure the CGCF using the redshift-space coordinates, with the BCGs at the origin by definition (by redshift space, we mean including all terms in equation 17 here). This should resemble the observed version of the CGCF, e.g. those measured in Wojtak et al. (2011), Sadeh et al. (2015) and Jimeno et al. (2015), except that we have turned the velocity differences into LOS distances to follow the convention in RSD. Like the GRedshift effect, all the second-order terms will induce front-back asymmetry in the CGCF.

(iii) To quantify the front-back asymmetries, we fit the CGCF at different projected distances σ from the centre with a Gaussian function and locate the shifts of the peaks in the CGCF. This provides a measurement for the amplitudes of the signal as a function of σ .

Our procedure of measuring the CGCF and quantifying the amplitudes of the front-back asymmetries is equivalent to that adopted in Wojtak et al. (2011), Sadeh et al. (2015) and Jimeno et al. (2015). Our simulation results therefore quantify directly the impact of the second-order terms in equation (17) on those measurements. Note, however, that quantifying the GRedshift signal in terms of the shifts of the peaks in the CGCF is different from measuring the mean redshift difference (Bonvin et al. 2014), e.g. $\int_{z_{\min}}^{z_{\max}} (z_{\text{BCG}} - z)\xi(\sigma, z)dz$, where z_{\min} and z_{\max} are the lower and upper bound of integration around the cluster, and ξ is the CGCF. The latter may depend strongly on the symmetry near the integration boundaries, while the shifts of the peaks are unlikely to be affected by those. It is the latter that we are quantifying and that has been measured by Wojtak et al. (2011), Sadeh et al. (2015) and Jimeno et al. (2015).

6.1 GRedshift in real space

As an intermediate step and for the purpose of comparison, we first measure the impact of GRedshift in the CGCF in real space, i.e. without the perturbing effect of peculiar velocities. In the simulations, we turn off the velocity term in equation (17) so that the particles are displaced only by GRedshifts. The cluster-galaxy (or halo centre-particle) cross-correlation function (CGCF), $\xi(\sigma, \pi)$ shown in the top left-hand panel of Fig. 4, is computed using all particles. For this figure, the amplitudes of the Φ s have been artificially boosted by a factor of 100 to aid visualization. The cluster centres are redshifted most as they are at the bottom of the potential wells. The amount of GRedshift decreases as the distance relative to the cluster centre increases. Particles are therefore preferentially shifted towards the observer at large separations both in front of and behind the cluster centre. This effect decreases with projected distance, causing the inverted-candle-flame shape for the correlation function. Relative to the cluster centre, the rest of the particles are blueshifted with an amplitude that generally increases with projected distance, but with the exception of when the impact of neighbouring groups or clusters is significant.

The amplitudes of the GRedshifts are measured by locating the peak of the particle distribution function (PDF) at a given projected distance σ from the stacked cluster centre. The peaks are located by fitting the PDFs with a Gaussian plus a constant:

$$f(y) = A + B \exp[-(y + \hat{\Phi}^{\text{obs}})^2/C], \quad (21)$$

where A , B , C and $\hat{\Phi}^{\text{obs}}$ are free parameters, and $\hat{\Phi}^{\text{obs}}$ is the parameter of interest. The middle left-hand panel shows examples of the measured PDFs (solid) and the best-fitting results (dotted). The poor agreement between the solid and the dotted lines away from the peaks does not affect our results, as we are only interested in the locations of peaks.

The dashed curve in the bottom left-hand panel is the recovered amplitude of Φ^{obs} . A bump at $\sim 2 \text{ Mpc } h^{-1}$ caused by neighbouring clusters is clearly seen. This is consistent with what we found when measuring the potential profiles directly (in real space) as shown in Fig. 3.

Note that in practice, with realistic amplitudes of the Φ s, it is nontrivial to recover the asymmetry of the correlation function purely due to GRedshift. Sample variance will strongly affect the

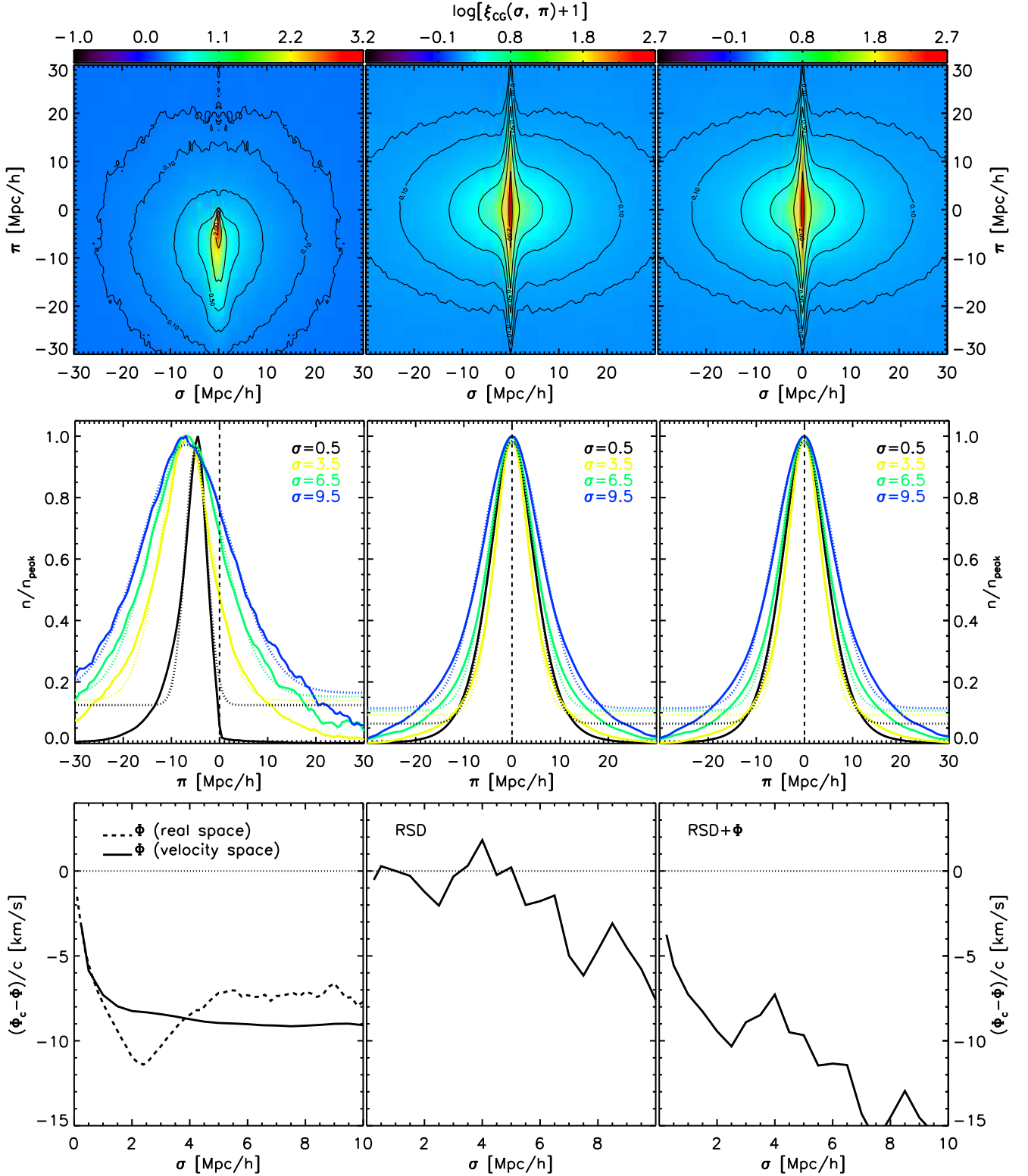


Figure 4. Top panels: the cluster-mass correlation functions in GRedshift space (left-hand panel), in peculiar velocity space (RSD) (middle panels) and the sum of the two (right-hand panel) for haloes with the mass $M > 5 \times 10^{13} M_{\odot} h^{-1}$. In the left-hand panel, no peculiar velocities are added and the GRedshift signal has been amplified by 100 times for better visualization. In the right-hand panel, the GRedshift distortion is much smaller than that from peculiar velocities and it is difficult to see the difference it produces relative to the middle panel. Middle panels: examples of histograms of the particle distributions (from the top panels) along the LOS direction, π , at different projected distances, σ , from the cluster centres. Dotted lines are the best-fitting models (equation 21) to the solid lines. The offsets of the fitted peaks from the centre are interpreted as the GRedshift signal. Bottom panels: the best-fitting values for the offsets from the centre shown in the middle row. Subtracting the curve on the right from the one in the middle yields the solid curve on the left, which is the velocity space version of the GRedshift signal. The dashed curve on the left is the real space version. The non-zero values in the middle panel are due to sample variance.

Table 1. Estimated signal to noise for the GRedshift signal from stacked clusters from the Millennium simulation. The Φ values are taken from the minima of the stacked profiles.

M_{\min} ($M_{\odot} h^{-1}$)	σ_v (km s^{-1})	N_{halo}	Φ (km s^{-1})	S/N for 1/8 ($\text{Gpc } h^{-1}$) ³	S/N for 1 ($\text{Gpc } h^{-1}$) ³
1×10^{13}	240	35 300	3	2.3	6.6
5×10^{13}	400	5283	10	1.8	5.1
1.6×10^{14}	600	1000	15	0.7	2.1
3.5×10^{14}	800	180	22	0.3	1.0

measurement. Although in principle, it can be beaten down by using very large samples, we are limited by the size of the simulation. It is therefore necessary to employ another technique to suppress the variance. Boosting Φ helps to illustrate the effect, but we think that this method is not ideal as sample variance is not completely eliminated. Instead, in the next section, we adopt another method to eliminate sample variance. This new method is found to be robust regardless of the amplitude of the signal.

6.2 GRedshift in velocity space

6.2.1 Signal to noise

In reality, measurements are made in redshift space including Φ , etc., but from equation (17), it is obvious that the peculiar velocity term is the most dominant and so to a good approximation, all other quantities are measured in velocity space. The GRedshift signal is of the order of 10 km s^{-1} for haloes with masses above $5 \times 10^{13} M_{\odot} h^{-1}$. This is at least one order of magnitude smaller than the peculiar velocity dispersion of clusters σ_v . Suppose we have $\sigma_v = 400 \text{ km s}^{-1}$, and we want to achieve a 3σ detection of the GRedshift, then we will need to stack of the order of 14 000 clusters. We list the estimated signal to noise of the GRedshift signal for a few cluster samples with different velocity dispersions in Table 1. The signal to noise increases with decreasing σ_v or the minimum halo mass M_{\min} . A 5σ detection can be achieved with a group sample of $M > 5 \times 10^{13} M_{\odot} h^{-1}$ or $\sigma_v > 400 \text{ km s}^{-1}$ in a volume of 1 ($\text{Gpc } h^{-1}$)³. Within the dynamical range of haloes from Table 1, Φ versus σ_v can be fitted by a linear function $\Phi/c = (\sigma_v/30 - 5) \text{ km s}^{-1}$. Or $\Phi/c = [18(M_{\min}/10^{13})^{1/3} - 5] \text{ km s}^{-1}$.

In the Millennium simulation, we have 5283 clusters with $M > 5 \times 10^{13} M_{\odot} h^{-1}$. This is only one-third of the halo number needed achieve a 3σ detection. To increase the sample without using a larger box-size simulation, we view the simulation along its three principal axes. This effectively increases the number of haloes by a factor of 3.

6.2.2 Redshift-space distortions only

To highlight the impact of sample variance, we show in the middle column of Fig. 4 results without GRedshift, and with peculiar velocity distortions only. This is the case of conventional RSD. In principle, without sample variance, no asymmetry along the LOS is expected. When we follow the same procedure by using equation (21) to fit for Φ^{obs} , it should be zero. However, from the bottom-middle panel, we see that the measured Φ^{obs} fluctuates around zero at the level of a few km s^{-1} . The expected dispersion of Φ^{obs} estimated from the velocity dispersion and size of this sample is approximately 3 km s^{-1} . The amplitude of fluctuations are consistent with sample variance.

6.2.3 GRedshift + RSD

Finally, in the right-hand column of Fig. 4, we have both the GRedshifts and peculiar velocities turned on. Both the correlation functions and the PDFs look essentially identical to the case of RSD only as the additional GRedshifts are much smaller than the peculiar velocity distortions. For the best-fitting Φ^{obs} , even though they are noisy, we see offsets of about 10 km s^{-1} when comparing the bottom right-hand panel with the bottom-middle panel. The difference between them, as shown by the solid curve in the bottom left-hand panel, is *roughly* consistent with the dashed curve shown in the same panel, which is the GRedshift signal recovered in real space.

In observations, the recovered GRedshift signal should be something like the bottom right-hand panel. It is affected by sample variance. The effect of sample variance can overwhelm the signal if the sample is too small. From simulations, we can effectively reduce sample variance by subtracting from the case of GRedshift + RSD the result of RSD only. This yields the GRedshift signal free from sample variance. We find that this method is robust regardless of the amplitudes of the GRedshift signal. The result is shown by the solid curve in the bottom left-hand panel.

To double check for the robustness of this method, we test using another technique to eliminate sample variance. We view each cluster from two opposite directions and stack them together before performing the fitting. This guarantees that each stacked cluster is perfectly symmetric along the LOS in velocity space. The pure GRedshift signal can then be recovered. We find that the recovered GRedshift signals from these two methods are consistent with each other. We will later apply them to measure the other quantities on the right-hand side of equation (17).

Quantitatively, the recovered GRedshift signal in velocity space (solid curve in the bottom left-hand panel of Fig. 4) is found to be different to the real space version (dashed curve). This indicates the strong influence of the peculiar velocity on the observed GRedshift signal. Two more examples of this comparison are shown in Fig. 5, where the pure GRedshift signals are shown by the red curves. The real space GRedshift signal recovered from the CGCF is consistent with the measurements shown in Fig. 3. The troughs at $\sim 2 \text{ Mpc } h^{-1}$ indicate that the impact of neighbours is again important. In velocity space, however, those troughs no longer exist and the GRedshift profiles are very different from their real space counterparts.

The difference of Φ^{obs} in real and redshift space is not surprising. The observed redshift of particles or galaxies with large velocities relative to the cluster centre will appear far away from their original positions in the cluster system. This will alter their distributions along the LOS, shifting the peaks of the PDFs relative to the cluster centre. The observed GRedshift signal in velocity space is therefore different from its original real space version. Note that due to the domination of the velocity dispersion over other effects of interest, all the other terms on the right-hand side of equation (17) will also be significantly altered in velocity space. In order to match observations, it is therefore important to make model predictions of this kind in velocity space.

6.3 Other second-order terms in the past light cone

We now discuss in some detail the various terms occurring in equation (17). The reader who is not interested in such details may skip to the conclusions.

Using the same technique as the previous section, we quantify the effect of the other terms on the right-hand side of equation (17). In light of the strong impact of peculiar velocities on the predicted

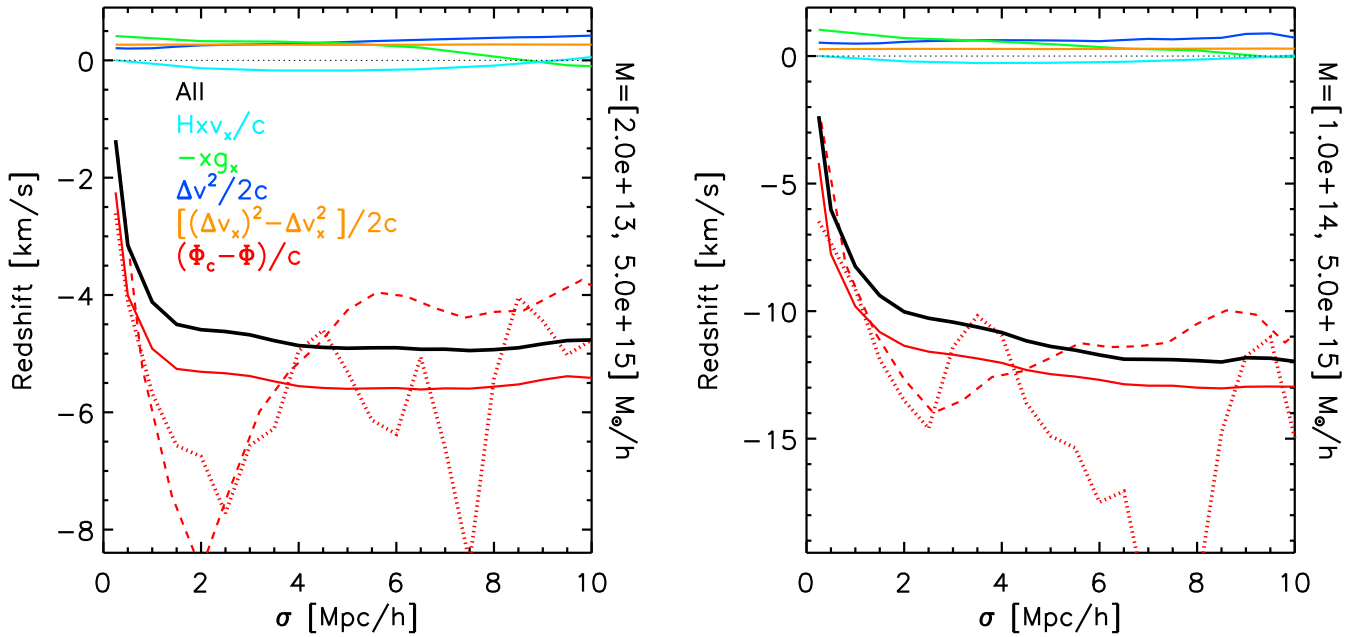


Figure 5. The dependence of the redshift offset of the peak of cluster–mass cross-correlation function on transverse separation, σ . The two panels are for different halo mass ranges as labelled on their right-hand axes. $(\Phi_c - \Phi)/c$ (red solid) is the idealized GRedshift signal in velocity space. The red dashed curves are their *real space* versions. They are comparable to the solid curves in Fig. 3. Blue and orange represents quantities related to the special relativistic effect. Green and cyan curves represent the other two terms arising from the effects of the past light cone. The black curves are the sum of all the terms. The actual measurements from simulations including sample variance are shown by the red dotted curves.

GRedshift signal shown previously, we show results only in velocity space for the other second-order terms, but we have also checked explicitly their real space counterparts to gain a better understanding of the physics.

(i) Δv^2 : Labelled as $\Delta v^2/2c$ in Fig. 5, the special relativistic correction term always produces a redshift and so is found to have the opposite sign to the GRedshift signal (blue curve), consistent with the results of (Zhao et al. 2013). However, the amplitude of the signal turns out somewhat smaller, i.e. at the sub- km s^{-1} level. When we examine its real space version, we find that there is a peak within the virial radius and its amplitude is approximately a factor of 2 larger than that in velocity space. This can be understood by the fact that particles having large peculiar velocities are displaced in v -space from their original locations. The v -space version therefore turns out to be smoother and have no obvious peak.

(ii) $[(\Delta v_x)^2 - \Delta v^2]/2c$: This can also be written as $(v_{xc}^2 - v_{xc}v_x)/c$, from which it can be seen that by definition, it vanishes at the position of the BCG. At a non-zero distance from the cluster centre, the second term should be very small when averaged over a large sample, leaving v_{xc}^2 as the dominant term. So this is effectively the special relativistic correction arising from the non-zero velocity dispersion of the BCG. From the orange curves in Fig. 5, we see that it is nearly a constant as expected.

(iii) Hxv_x/c : The term Hxv_x is shown by the cyan curves. It is the product of the radial Hubble flow with the LOS peculiar velocity. In the virialized region, Hx and v_x are uncorrelated because the peculiar velocities are random. In the outskirts of a cluster, they are anticorrelated because of infall, i.e. Hx is positive and v_x is negative. This remains the same until the peculiar velocities drop to zero at very large distances, where they are back to no correlation. Initially, one might expect that in real space, this induces negative redshifts (blueshifts) with respect to the cluster centre, the same as

the GRedshift signal. This is true for individual particles or galaxies, but what we find is that peak of the particle distribution is actually redshifted. However, when switched to velocity space, the sign of this term is reversed again to become a blueshift.

Fig. 6 shows an example to explain all these subtleties. Initially, the PDFs of particles along the LOS are symmetric about the centre in real space, as shown by dotted curves in the right-hand panel. When adding the term $200 Hxv_x$ (dashed curves), the PDFs are skewed and the peaks are shifted to the positive π direction, even though individual particles move in the negative π direction. This happens because of the joint effect of the amplitude of Hxv_x increasing and the amplitude of the PDF decreasing with increasing LOS distance. At large positive π , particles are shifted towards the centre, which increases the amplitude of the PDF near the centre. At large negative π , particles are shifted away from the centre, causing a decrease of the amplitude of the PDF. The consequence is that the peak of the PDF is shifted in the positive π direction. The shift is more pronounced at large σ (orange and red dashed curves) as the amplitudes of $200 Hxv_x$ is larger.

In velocity space (solid curves), it is noticeable for the orange curve that the peak of the PDF is shifted in the negative π direction. This is because the sign of the average velocity along the LOS v_x is flipped in velocity space at relatively small σ values. This is shown by the left-hand panel of Fig. 6; the LOS v_x is negative at all values of σ as expected from the infall motion of mass towards the stacked cluster centre. However, random particle velocities close to the cluster centre displace particles with positive velocity to positive π coordinate in velocity space and vice versa. This reverses the correlation between v_x and π . This only occurs at relatively small σ values, i.e. $\sigma \lesssim 9 h^{-1} \text{Mpc}$. Therefore, the effect of the term Hxv_x in velocity space is to cause blueshifts at $\sigma \lesssim 9 h^{-1} \text{Mpc}$, which is the same as the GRedshift effect, but it gives rise to redshifts at $\sigma \gtrsim 9 h^{-1} \text{Mpc}$. This is shown by the cyan line in Fig. 5.

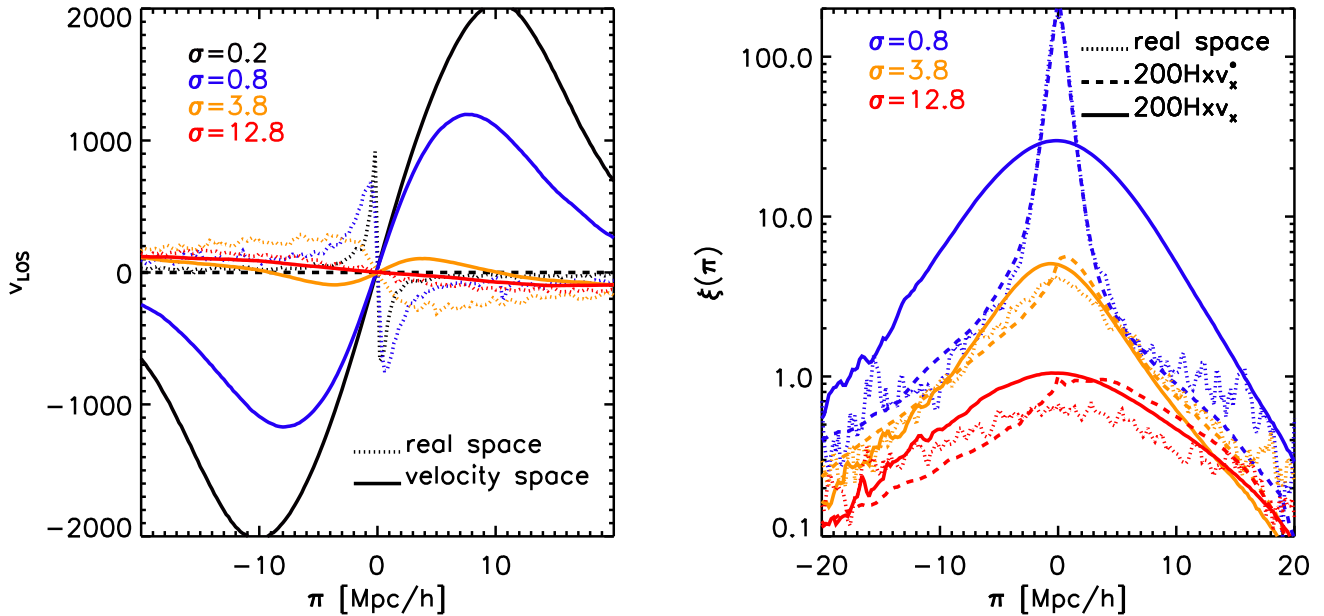


Figure 6. Left-hand panel: averaged LOS peculiar velocities of particles as a function of distance from the cluster centre in real space (dotted lines) and in velocity space (solid lines) at different projected distances indicated by the different colours. At small σ values, the sign of the LOS velocity in velocity space is flipped with respect to its real space version due to velocity dispersion and infall. Right-hand panel: distribution of dark matter particles along the LOS for three different cases: (1) real space (dotted lines); (2) with the LOS real space positions of particles perturbed by the second order term $200Hxv_x^*$, where the boost factor of 200 is used for better illustration (dashed); (3) the LOS velocity space positions of particles are perturbed by $200Hxv_x$. The peaks of the particle distribution for case (1) are expected to be at the centre. In case (2), they are shifted to the positive π direction (redshift) and the distribution is skewed due to the effect of the $200Hxv_x$ term. In case (3), at relatively small σ values (the orange and blue curve), the peaks are shifted to negative π (blueshift) due to the fact that the sign of the LOS velocity is flipped in velocity space, as shown in the left-hand panel. These results are for haloes within the mass range of $M > 10^{14} M_{\odot} h^{-1}$.

(iv) $-xg_x$: The green curves show the effects of the term $-xg_x$, minus the product of the LOS displacement and the LOS acceleration. We can understand it as arising from the change of the velocity of the galaxy with respect to the cluster centre during the interval of look-back time between the galaxy and cluster centre. It is defined to be zero at the cluster centre. At a non-zero projected distance from the cluster centre, x and g_x tend to be anticorrelated (for overdense systems), i.e. the acceleration will decrease (becoming less negative) with increasing distance from the cluster centre. With the negative sign, we expect each individual particle (or galaxy) to be redshifted (positive redshift) with respect to the cluster centre. Initially, one may expect that in real space, this term will have the opposite sign to that of the GRedshift effect. However, for the same reasons as those for the Hxv_x term, the peak of the CGCF is found to be shifted towards negative π side (far side) of the centre in real space, and the sign of the recovered $-xg_x$ term flips again at $\sigma < 9 \text{ Mpc } h^{-1}$ in velocity space due to the infall velocities and dispersion. Therefore, the sign of the measured signal for the $-xg_x$ term is redshift at $\sigma < 9 \text{ Mpc } h^{-1}$ and blueshift at larger projected distances, as shown by the green curves in Fig. 5.

The quadratic (in x) term that comes from the combination of the background GRedshift and Doppler effects is assumed to be removed by fitting the background density ramp to the LOS galaxy distribution around the cluster centre, as reasoned in Section 2. We therefore do not include it in our figure.

Finally, the contribution of all these terms to the overall redshift signal are shown by the brown curve in Fig. 5. They reduce the amplitude of the GRedshift signal (red solid curve) by approximately 0.5 and 1 km s^{-1} for the two halo samples presented in Fig. 5. This is relatively minor (as some of them cancel with each other) compared

to the other two systematics (the impact of neighbours and the combined effect of velocity space) identified earlier. With the effective volume of $3 \times [0.5(\text{Gpc } h^{-1})]^3$, the expected observed GRedshift signal is shown by the black solid curves. All the systematics are overwhelmed by sample variance, which is reflected by the strong fluctuations of the curves.

7 DISCUSSION AND CONCLUSIONS

We have explored how the modelling of the GRedshift signal from stacked clusters is affected by a variety of systematics.

(i) Since the GRedshift signal is a component on the observed redshift, we start by presenting the expression equation (17) for the observed redshift on the past light cone of an observer including relativistic corrections. It is relative to the centre of a cluster, and is expressed in terms of properties on surfaces of constant proper time. The effect of the second-order terms in this expression on the cluster-galaxy cross-correlation function (CGCF) are quantified using N -body simulations. We find that the GRedshift term causes the strongest asymmetry of the CGCF. The recovered GRedshift signal is biased high by approximately 0.5–1 km s^{-1} depending on the minimum halo mass due to neglecting the other second-order terms. This is relatively minor compared to the other two other systematics we have found.

(ii) The underlying gravitational potentials are usually deeper where there is a concentration of galaxies, which indicates a concentration of mass. The fact that observations of GRedshift are galaxy number-weighted causes the observed GRedshift signal to be biased low compared to models where volume weighting is assumed. This bias does not go away even if the stacked cluster is

perfectly spherically symmetric. The non-spherical distribution of galaxies in individual clusters and the complex cosmic web structures surrounding the cluster cause the bias to persist at nearly all scales of interest. This bias is stronger for lower mass clusters as the chance of having more massive neighbouring structures is higher. A pronounced bump at approximately $2 \text{ Mpc } h^{-1}$ from the cluster centre is expected for the observed GRedshift profile due to this bias. However, the bump tends to be flattened in velocity space.

(iii) Peculiar velocities of galaxies are the most dominant feature in the CGCF. The measurement of the GRedshift signal is, in essence, conducted in velocity space. It is strongly influenced by peculiar velocities since the observed galaxies are shifted from their original locations, e.g. galaxies at the bottom of the potential may appear far away from the cluster centre due to velocity-space distortions. This tends to flatten the bump of the GRedshift profile caused by the impact of neighbouring structures as mentioned in the previous point. It also affects the predictions for all the other second-order terms in equation (17).

(iv) We find that the CGCF along the LOS associated with the GRedshift signal is highly non-Gaussian. Therefore, extracting the signal by using a Gaussian function to fit for the peak positions of the CGCF as done in Wojtak et al. (2011), Sadeh et al. (2015) and Jimeno et al. (2015) may not be the optimal. There may be room for improvement in future analysis of this kind. It is also worth noting that extracting the shifts of the peaks in the CGCF is different from measuring the mean redshifts of galaxies with respect to the BCG. The latter may depend more strongly on the boundaries of integration along the LOS. The box-size of the simulation we use in this study is relatively small. The methods we have developed allow us to extract the relatively weak signal free from sampling variance. Simulations with larger box-size will be needed to study the noise properties.

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